

Exercises Astronomical Observing Techniques, Set 3

Exercise 1

- Calculate the theoretical resolution (in arcsec) of a 3.6m telescope observing at a wavelength of 500 nm.
- The Fried parameter (r_0) of the atmosphere is 20 cm at a wavelength of 500 nm. Calculate resolution of the telescope (in arcsec) looking through the atmosphere.
- What is resolution (in arcsec) at $2 \mu\text{m}$ using the information in 1b?
- A $2\text{k} \times 2\text{k}$ CCD camera is placed in Prime Focus operating at 500 nm. The telescope has a focal length (f) of 10 m, we want to have 4 pixels per resolution element using the atmospheric conditions described above (1b). Calculate the pixel size and the resulting Field of View (FOV).

Exercise 2

Discuss the advantages/disadvantages of the Cassegrain over the Nasmyth focus.

Exercise 3

Discuss the advantages/disadvantages of space observations vs ground observations.

101.6	110.9	105.7	116.9	94.7	104.8
106.7	93.9	103.4	94.5	89.1	88.2
110.8	111	113.3	94.2	85.2	101.7
117.3	100.8	113.1	130.2	106.5	101.2
108.1	105.7	94.6	90.1	106.2	105.7
101.6	95.3	99.9	83.1	70.7	100.6

Figure 1: Pixel values

Exercise 4

Consider a perfect (i.e., no detector noise) but small, 6×6 pixel detector array, which is exposed to a uniform sky background that will produce a constant flux level plus the associated photon shot noise. In addition to the sky background we are observing a very faint star. For simplicity, we assume that

- the flux from the star will be uniformly spread across four pixels
- the star will illuminate the central four pixels of the array

- the signal from the star has no associated shot noise.

The measured pixel values of the array are given in Figure 1.

a) Calculate the mean and the standard deviation of the pixel values of the array. Since we know that the central four pixels contain the source we exclude them from the statistics. What is the relative noise ($1\sigma/\text{mean}$) in percent?

b) Now we consider if we can claim a detection of a source on a pixel-by-pixel basis. We know that the four central pixels see the source flux. What are the significances of detection (in standard deviations) for each of the four pixels?

c) In order to improve the S/N we now re-bin the 6x6 onto a 3x3 array. In other words, we combine (coadd) the values from 2x2 neighboring pixels. What are the pixel values now? Draw a 3x3 table.

d) Calculate the mean and the standard deviation of the re-binned array. Again, exclude the central pixel, which contains the source, from the statistics. What is the relative noise in percent? How does it compare with the previous "single pixel" S/N?

e) What is the significance of detection of the star in the central pixel? Is the level of confidence high enough to claim a reliable detection?